IAU Colloquium 164: Radio Emission from Galactic and Extragalactic Compact Sources
ASP Conference Series, Vol. 144, 1998
J. A. Zensus, G. B. Taylor, & J. M. Wrobel (eds.)

The Identification of EGRET Sources with Flat-Spectrum Radio Sources

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Abstract. We provide a table of 42 blazars which we expect to be robust identifications of EGRET sources. The table includes the lensed source, PKS 1830−211 which we identify with a previously unidentified EGRET source.

1. Introduction

A careful analysis of the identifications of EGRET sources with radio sources is in order because the source of the γ-ray emission is not well located. An analysis of all EGRET sources off the Galactic plane in the 2nd EGRET catalog and its extension (Thompson et al. 1995, 1996) has been published by Mattox et al. (1997).

2. The Catalog

The catalog of 42 sources for which the confidence of the identification exceeds $\sim$90% is given in the table. Column 1 is the epoch 1950 position name of the radio source. Column 2 gives the 5 GHz flux density and column 3 the cm radio spectral index, $S(\nu) \propto \nu^{\alpha}$. Column 4 gives the correlated VLBI flux density at 2.29 GHz from Preston et al. (1985). Column 5 gives the probability that the identification is corrected based on EGRET data and radio data. In some cases increased confidence results from variability of the γ-ray flux or a high VLBI flux density. See Mattox et al. (1997) for details. Column 6 gives the peak flux detected by EGRET, and column 7 gives the EGRET photon spectral index, $dF(E) \propto E^{-\gamma} dE$. During the analysis done to produce this catalog, Mattox et al. (1997) found that:

1) It appears that only the blazar class of AGN is being detected.

2) The EGRET sources have higher VLBI flux densities than the parent population of flat radio spectrum sources.

3) The peak γ-ray flux of a blazar is correlated with its average 5 GHz radio flux. An even better correlation is seen between γ-ray flux and VLBI flux.

References

Mattox, J. R., et al. 1997. *ApJ*, **481**, 95–115.
Preston, R. A., et al. 1985. *AJ*, **90**, 1599–1641.
Thompson, D. J., et al. 1995. *ApJS*, **101**, 259–286.
Thompson, D. J., et al. 1996. *ApJS*, **107**, 227–237.

Name B1950	S_5	α	$S_{\rm VLBI}$	$p(id\|r)$	Peak Flux 10^{-6}cm^{-2}s^{-1} E>100 MeV	γ-Ray Spectral Index
0202+149	2714	−0.4	600±100	0.994	0.26±0.06	2.5±0.1
0208−512	3198	−0.2	2230±110	0.996	1.10±0.07	1.7±0.1
0234+285	3356	0.3	1600±100	0.95	0.16±0.04	2.7±0.3
0235+164	1955	−0.1	1800±200	0.95	0.82±0.09	2.0±0.2
0336−019	3014	0.3	1400±100	0.997	4±1	2.4±0.3
0420−014	6992	1.1	610±50	0.9995	0.50±0.10	1.9±0.3
0440−003	1084	0.6	820±40	0.91	0.84±0.12	1.8±0.2
0458−020	3317	1.0	920±80	0.96	0.31±0.09	2.5±0.4
0521−365	8180	−0.7	990±90	0.9989	0.21±0.06	2.2±0.4
0528+134	2978	0.3	500±80	0.999	2.95±0.33	2.6±0.1
0537−441	4805	0.4	2030±90	0.998	0.36±0.09	2.0±0.2
0716+714	810	0.0	600±50	0.78	0.53±0.13	1.9±0.2
0735+178	2146	0.1	600±100	0.96	0.30±0.12	2.5±0.5
0827+243	670	0.0	800±100	0.34	0.26±0.06	2.2±0.4
0829+046	2105	0.7	500±100	0.86	0.20±0.06	2.5±0.5
0836+710	2436	−0.4	370±30	0.79	0.45±0.11	2.4±0.2
0954+556	2260	−0.2	< 100	0.989	0.11±0.04	1.8±0.3
0954+658	1419	0.6	430±30	0.992	0.14±0.04	1.7±0.2
1101+384	712	−0.1	320±40	0.70	0.23±0.08	1.7±0.2
1127−145	4209	0.	1490±70	0.90	0.93±0.23	2.7±0.4
1156+295	1542	−0.1	430±50	0.94	2.29±0.55	2.0±0.5
1219+285	968	0.	400±40	0.74	0.28±0.09	1.3±0.4
1222+216	1261	−0.4	330±40	0.40	0.83±0.20	1.9±0.4
1226+023	44940	−0.1	1510±70	0.9995	0.24±0.05	2.4±0.1
1253−055	11192	−0.1	3700±300	0.999991	4.5±0.6	1.9±0.1
1406−076	1080	0.2	640±50	0.86	1.44±0.26	2.0±0.1
1424−418	2597	0.0	370±40	0.98	0.55±0.17	2.6±0.4
1510−089	3000	0.0	2300±130	0.98	0.48±0.18	2.6±0.4
1606+106	1688	0.0	490±50	0.95	0.60±0.13	2.2±0.3
1611+343	2483	−0.1	1200±100	0.95	0.55±0.13	2.0±0.2
1622−253	3449	0.7	150±20	0.998	0.43±0.07	2.3±0.2
1622−297	1920	0.1	2000±200	0.996	17±3	1.9±0.2
1633+382	3198	0.4	1300±100	0.9996	1.8±0.3	1.9±0.1
1730−130	6991	0.8	1420±60	0.988	1.37±0.43	2.4±0.3
1739+522	1134	−0.4	650±70	0.61	0.54±0.11	2.2±0.4
1830−211	7920	−0.3		0.98	0.70±0.19	2.7±0.3
1908−201	2053	0.	460±50	0.95	0.29±0.10	2.5±0.2
2052−474	2026	0.	380±40	0.90	0.25±0.08	2.4±0.4
2200+420	3571	−0.2	720±60	0.993	0.40±0.12	2.2±0.4
2209+236	1194	0.7	550±50	0.81	0.13±0.04	2.8±0.5
2230+114	3765	−0.5	600±100	0.998	0.53±0.22	2.6±0.2
2251+158	15859	0.1	3400±300	0.9998	1.32±0.21	2.2±0.1